

# Tired Light and Type Ia Supernovae Observations

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## Abstract

The phenomenon known as “supernovae timescale stretching” does not rule out “tired light” theory but actually can be explained within redetermined “tired light” paradigm. Using Type Ia supernova spectra template I have conducted computer simulation study of light propagation from supernova to observer based on the “tired light” assumptions. The results of simulation for UBVRIYJHK bands “at rest frame” do not contradict to astronomical observations related to light curve broadening and spectra aging. The comparative analysis shows advantages of “tired light” over “accelerated Universe expansion” model.

## 1. Introduction

In 2001 group of cosmologists from Supernova Cosmology Project lead by Gerson Goldhaber published article “Timescale Stretch Parameterization of Type Ia Supernova B-band Light Curves” [1]. While analyzing explosions of Type Ia supernovae they discovered dependency of the width of B-band light curves on supernova redshift. They found that the light curve width is proportional to factor  $(1+Z)$ , where  $Z$  is redshift. In the article they concluded that this dependency “*provides independent evidence for cosmological expansion as the explanation of redshifts.*” The important part of the work is the proof of universality of discovered phenomenon of timescale stretching. The authors rule out “tired light” theories as they “*would not yield this slowing of the light curves*”.

In 2008 group of cosmologists lead by Stéphane Blondin suggested another approach for measuring timescale stretching of supernova explosion. In their work “Time Dilation in Type Ia Supernova Spectra at High Redshift” [2], they suggested to measure timescale stretching by spectra aging. They found that spectra of low redshift supernovae changes in very consistent way. By comparing the rate of spectra aging for low and high redshift supernovae it is possible to measure timescale stretching factor. They found that aging rate is proportional to  $1/(1+Z)$  factor, which conforms the results of Goldhaber group and supports “*homogeneous and isotropic expanding universe*”. Blondin et al. also rule out “tired light” hypothesis as “*this hypothesis does not predict a time dilation effect*”.

In my work I show that “tired light” hypothesis does not contradict to astronomical observations related to light curve broadening and spectra aging of Type Ia supernovae. Moreover, these observations can be explained within “tired light” paradigm.

## 2. Redetermined “tired light” paradigm

Term “tired light” covers a class of theories that explain redshift by photon energy depletion effect. These theories differ by physical mechanism of energy depletion. The suggested mechanisms include the following: Compton Effect, “gravitational drugging” (Zwicky 1929), “some unknown effect” (Hubble, Tolman 1935), traversing a radiation field (Finlay-Freundlich, 1954), curvature pressure (Crawford 1999).

Karim Khaidarov suggested redshift mechanism based on the classical idea of interstellar space filled with universally present media – aether that is the carrier of electromagnetic waves. In this media waves experience “*constant fading, just like*

waves in the usual isotropic physical media: solids, liquids, gases.” (Khaidarov 2003). In his work Khaidarov explains the dependency of supernova magnitudes on redshift presented by Perlmutter et al. (1998) from the positions of aether physics.

The energy which is lost by light gets spent on heating aether (Khaidarov 2005). The temperature of aether in interstellar space is also known as CMB temperature and equals to 2.7K. Aether is a physical media which is attributed by physical properties such as pressure, density, temperature, elasticity, viscosity. The properties of aether define the speed of electromagnetic waves. These physical properties are not constant, they may fluctuate; therefore the speed of light in vacuum is not a fundamental constant. It would be odd to suggest that physical waves travelling on the same path for billion years will have exactly the same travelling time, second to second. The travelling time of photons issued from the same source and observed at the same destination varies. Based on the Central Limit Theorem one may suggest that the travelling time has Gauss distribution. The idea of photon travelling time having Gauss distribution was suggested by Alexander Chepik (2003). The longer travelling path, the higher is the standard deviation of photons travelling time. Higher standard deviation leads to longer observed duration of supernova explosion. In my study I express the standard deviation in days and assume that zero redshift corresponds to zero standard deviation. Another assumption is that the standard deviation is the same for all wavelengths, even though it is possible to suggest that the standard deviation may depend on wavelength. The latter is the subject for a future research.

### 3. Computer simulation of light propagation

To test the redetermined “tired light” hypothesis I have developed a computer program that simulates propagation of photons from supernova to observer. The simulation can be executed two modes: “tired light” and “standard cosmology”. The program is implemented in C# programming language. It is available for download as open source freeware (Holushko 2012).

For input data I use spectral template of Type Ia supernovae with zero redshift maintained by Eric Hsiao (Hsiao et al. 2007). This template is based on more than 1000 spectral measurements and contains data for each day (epoch) in the range from -20 to 84 epochs (days) of explosion. The data for each day contains spectral flux density characteristics in the range from 1000 to 25000 Angstrom binned to 10 Angstrom. Another portion of input data are filter transmission curves for UBVRIYJHK bands. I apply UBVRi transmission data of Johnson-Bessel filters used in MONET project (MONET 2011). Transmission data for YJHK filters was obtained from LBT observatory (LBT 2011).

The photons get emitted with interval of 0.01 day. On each emission the number of photons for each spectral bin is determined by formula

$$N_{\lambda}(t) = k \cdot I_{\lambda}(t) \cdot \lambda \quad (1)$$

$I_{\lambda}(t)$  - Spectral flux density for certain wavelength and moment of time in supernova reference frame. It is calculated on basis of spectral template using linear approximation.

$\lambda$  - Wavelength

$t$  - Moment of time in supernova reference frame, on the beginning of explosion  $t = 0$

$k$  - Arbitrary coefficient

The number of photons is a factitious value. The coefficient  $k$  is chosen in a way that provides number of photons large enough to ensure statistical quality and low enough to have reasonable simulation running time. The energy possessed by each photon on emission is calculated by formula (2).

$$e_{\lambda}(t) = \frac{I_{\lambda}(t) \cdot 1s \cdot 10\text{\AA} \cdot 1\text{m}^2}{N_{\lambda}(t)} \quad (2)$$

The value of  $e_{\lambda}(t)$  is also a factitious value which is required solely for reconstruction of spectrum at observer site.

For each emitted photon the program calculates the arrival moment of time  $t'$  in the observer reference frame. For “standard cosmology” mode the arrival moment of time is calculated by formula (3).

$$t' = t \cdot (1 + Z) + T_0 \quad (3)$$

$Z$  - Redshift

$T_0$  - Traveling time of the photons emitted at the beginning of explosion.

For “tired light” mode the arrival moment is calculated as following.

$$t' = t + Gauss(T, \delta) \quad (4)$$

*Gauss* - Random number generator with Gauss distribution

$T$  - Mean of photons travelling time

$\delta$  - Standard deviation of photons travelling time entered as input parameter

Values of  $T_0$  and  $T$  are chosen arbitrary. The only consideration is that they should be essentially larger than any observed duration of explosion. At observer site for each received photon the values of wavelength, energy and arrival time are known; therefore it is possible to reconstruct spectra of received light with or without redshift. The non-shifted spectra can be determined using formula (3). Hereinafter I use term “non-shifted” instead of “at rest” because in “tired light” model both supernova and observer are at rest.

$$I_\lambda(t') = \frac{\sum e_\lambda(t')}{1s \cdot 10\text{\AA} \cdot 1\text{m}^2} \quad (5)$$

A non-shifted light curve in terms of flux at certain band is determined as following.

$$F(t') = 1s \cdot \sum_\lambda [I_\lambda(t') \cdot r_\lambda \cdot 10\text{\AA}] \quad (6)$$

$r_\lambda$  - Filter transmission coefficient for specific wavelength

The non-shifted light curve in terms of magnitude is determined as following.

$$m(t') = -2.5 \cdot \text{Log}_{10} \left( \frac{F_{B\text{max}}}{F(t')} \right) \quad (7)$$

$F_{B\text{max}}$  - Maximum flux in B-band

The simulation output data includes spectra, flux light curves and magnitude light curves that are mapped to epochs in a way that epoch 0 corresponds to the moment of  $t'$  when flux in B-band has maximum value.

### 3. Light curve broadening in UBVR bands

The fact that spectral template of Eric Hsiao ends on 85 day since B-maximum imposes certain limitations on simulation results. The problem is that the photons what would be issued after 85<sup>th</sup> day but with travelling time less than the mean do not

contribute to simulation results. To cut off incorrect data I use the rule of “three sigma”. For example, if the standard deviation of travelling time  $\delta = 10$  days, then the results after 55 day since B-maximum should be ignored.

The magnitude light curves (non-shifted) for various standard deviations in UBVR bands are shown on Figure 1 with some magnitude offset for better visibility.

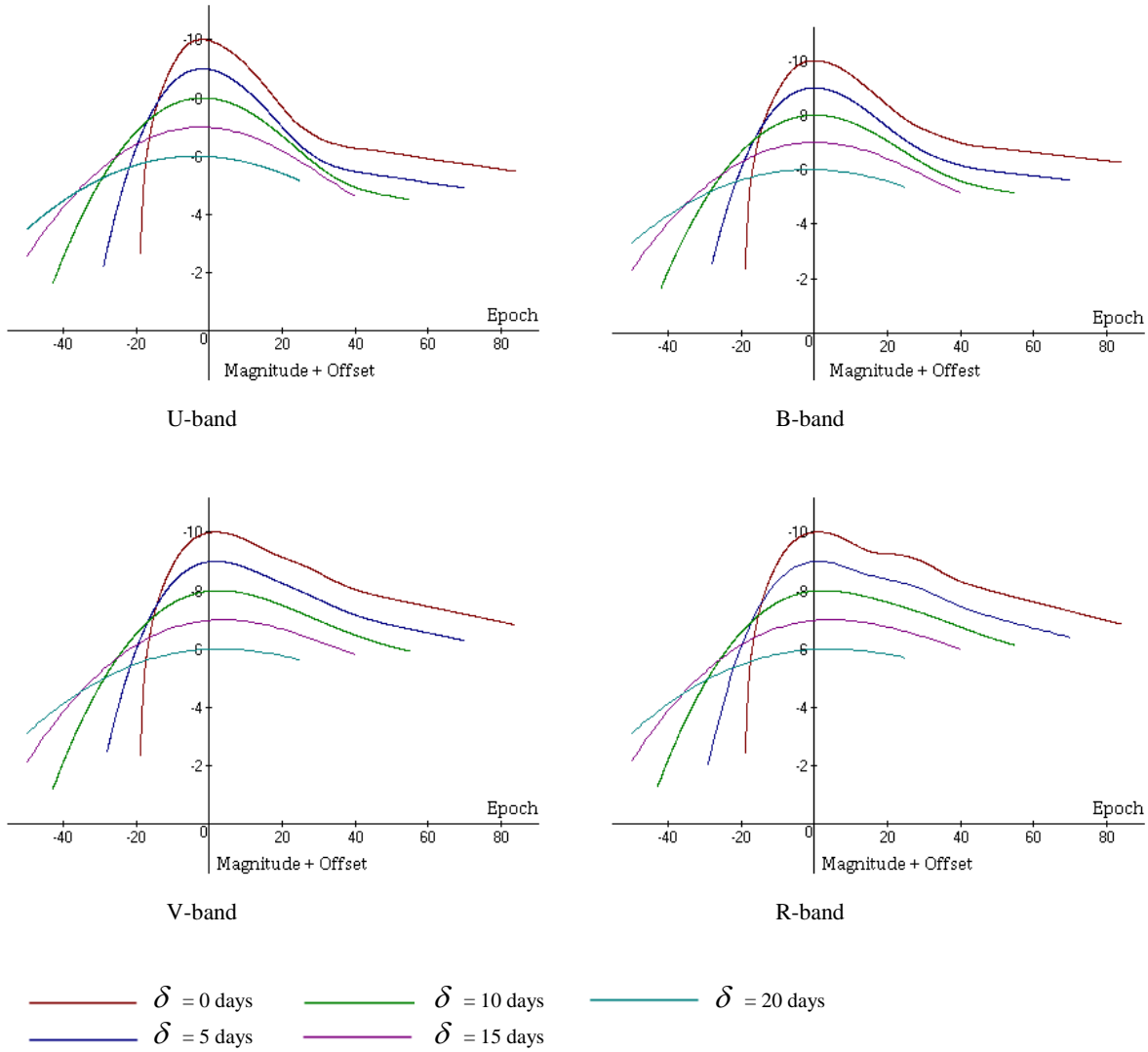


Figure 1  
The magnitude light curves in non-shifted UBVR bands

The width-factor  $w$  of timescale stretching was determined as a ratio between light curve widths with non-zero and zero standard deviations. Zero standard deviation corresponds to zero redshift. The measurements were taken for rising and falling fragments of the light curve as well as for the whole curve at certain magnitude offsets ( $\Delta m$ ) from maximum. The summary of measurements of light curve broadening in UBVR bands is shown in Table 1.

Table 1 - Measurements of light curve broadening in UBVR bands

Fragment	$\Delta m$	U		B		V		R	
		Width	$w$	Width	$w$	Width	$w$	Width	$w$
$\sigma = 0$ days, $Z=0$									
Rise	0.5	7.2	1	7	1	8.5	1	8	1

Rise	1	9.5	1	9.5	1	11.5	1	10.8	1
Rise	2	12.8	1	13.1	1	15	1	14.4	1
Rise	3	14.9	1	15.4	1	17.4	1	16.7	1
Fall	0.5	7.6	1	9.6	1	11.5	1	11.3	1
Fall	1	12	1	14.4	1	21	1	28.5	1
Fall	2	18.8	1	23.2	1	38.5	1	47	1
Fall	3	26.2	1	39.8	1	75.5	1	78.4	1
Whole	0.5	14.8	1	16.6	1	20	1	19.3	1
Whole	1	21.5	1	23.9	1	32.5	1	39.3	1
Whole	2	31.6	1	36.3	1	53.5	1	61.4	1
Whole	3	41.1	1	55.2	1	92.9	1	95.1	1
$\sigma = 5 \text{ days, } Z=0$									
Rise	0.5	8.3	1.15	9.0	1.29	10.1	1.19	9.6	1.20
Rise	1	11.4	1.20	12.4	1.31	13.8	1.20	13.1	1.21
Rise	2	15.8	1.23	16.9	1.29	18.6	1.24	19	1.32
Rise	3	18.9	1.27	20.2	1.31	22.0	1.26	21.4	1.28
Fall	0.5	9.9	1.30	10.6	1.10	13.4	1.17	14.9	1.32
Fall	1	14.4	1.20	15.8	1.10	22.6	1.08	29.2	1.02
Fall	2	21.9	1.16	25.6	1.10	42.2	1.10	50.6	1.08
Fall	3	30.7	1.17	46.0	1.16	-	-	-	-
Whole	0.5	18.2	1.23	19.6	1.18	23.5	1.18	24.5	1.27
Whole	1	25.8	1.20	28.2	1.18	36.4	1.12	42.3	1.08
Whole	2	37.7	1.19	42.5	1.17	60.8	1.14	69.6	1.13
Whole	3	49.6	1.21	66.2	1.20	-	-	-	-
$\sigma = 10 \text{ days, } Z=0$									
Rise	0.5	12	1.67	12.7	1.81	14	1.65	14.3	1.79
Rise	1	16.7	1.76	17.4	1.83	19.3	1.68	19.8	1.83
Rise	2	23.4	1.83	24.4	1.86	28.5	1.90	26.8	1.86
Rise	3	28.6	1.92	29.6	1.92	31.8	1.83	32.2	1.93
Fall	0.5	13.2	1.74	14.5	1.51	17	1.48	19.8	1.75
Fall	1	18.8	1.57	20.8	1.44	26.5	1.26	31.6	1.11
Fall	2	28.5	1.52	33.3	1.44	49.3	1.28	-	-
Fall	3	40.7	1.55	-	-	-	-	-	-
Whole	0.5	25.2	1.70	27.2	1.64	31	1.55	34.1	1.77
Whole	1	35.5	1.65	38.2	1.60	45.8	1.41	51.4	1.31
Whole	2	51.9	1.64	57.7	1.59	77.8	1.45	-	-
Whole	3	69.3	1.69	-	-	-	-	-	-
$\sigma = 15 \text{ days, } Z=0$									
Rise	0.5	16.5	2.29	16.70	2.39	17.60	2.07	18.70	2.34
Rise	1	23.2	2.44	23.50	2.47	24.70	2.15	26.00	2.41
Rise	2	32.6	2.55	33.10	2.53	34.50	2.30	35.70	2.48
Rise	3	39.6	2.66	40.20	2.61	42.00	2.41	43.20	2.59
Fall	0.5	16.9	2.22	18.20	1.90	21.50	1.87	23.70	2.10
Fall	1	24.4	2.03	25.40	1.76	33.00	1.57	36.20	1.27
Fall	2	37	1.97	42.20	1.82	-	-	-	-
Whole	0.5	33.4	2.26	34.90	2.10	39.10	1.96	42.40	2.20
Whole	1	47.6	2.21	48.90	2.05	57.70	1.78	62.20	1.58
Whole	2	69.6	2.20	75.30	2.07	-	-	-	-
$\sigma = 20 \text{ days, } Z=0$									
Rise	0.5	19.6	2.72	22.7	3.24	22	2.59	23.30	2.91
Rise	1	28.2	2.97	31.3	3.29	30.9	2.69	32.40	3.00
Rise	2	40.3	3.15	43.6	3.33	43.2	2.88	45.20	3.14
Fall	0.5	23.2	3.05	21.7	2.26	27	2.35	27.30	2.42
Whole	0.5	42.8	2.89	44.4	2.67	49	2.45	50.60	2.62

The measurement results in graphical form are shown on Figure 2 in form of dependency of width-factor  $W$  on standard deviation of travelling time  $\delta$  for various fragments and magnitude offsets.

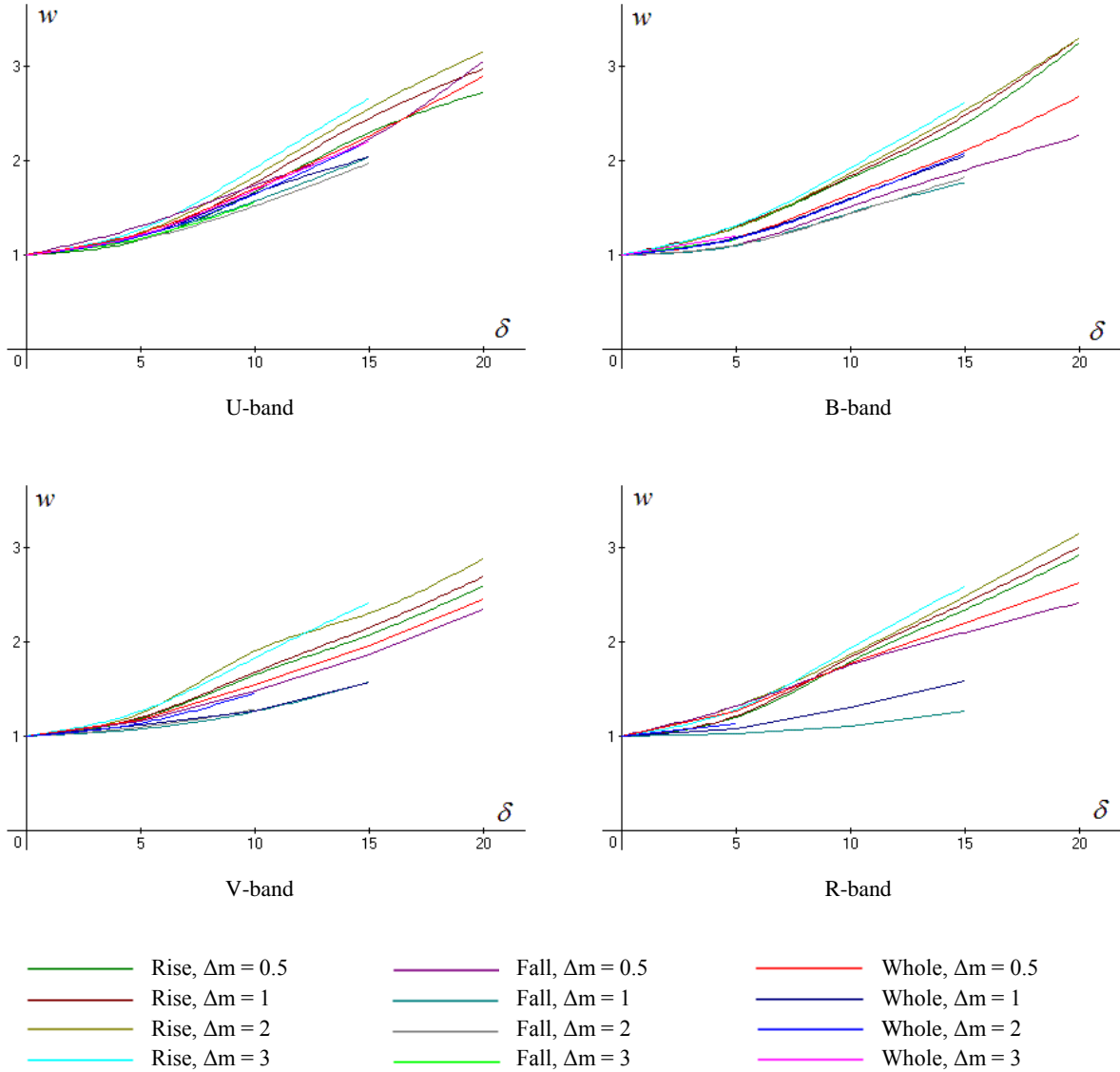


Figure 2

Dependency of width-factor on standard deviation of photons travelling time in non-shifted UBVR bands

The width-factor measurements presented in Table 1 and Figure 2 show that timescale stretching is not universal in "tired light" model. It varies depending on the band, magnitude offset and the fragment of a light curve. This contradicts the results reported by Goldhaber et al. (2001). Goldhaber group analyzed timescale stretching of Type Ia supernova light curves in B-band using the method of Composite Light Curve on the basis of photometric measurements of 12 low redshift supernovae from Calán/Tololo Survey and 35 high redshift supernovae reported by The Supernova Cosmology Project (SCP). They found it remarkable that the final composite light curve (Goldhaber et al. 2001, Figure 1e) appeared homogeneous. This result was presented as a proof of universality of timescale stretching. The intermediate transformations applied while building the composite light curve were based on the values of width-factor  $W$  and stretch-factor  $s$  measured by Perlmutter et al. (1997???) for the SCP high-redshift supernovae. The Perlmutter group measured light curve widths, width- and stretch-factors by fitting observed

photometric points to B-band light curve template and applying timescale stretching formula  $w=(I+Z)\cdot s$  universally. No wonder why Goldhaber's composite light curve looks so homogeneous. In simple words the logic of the Goldhaber group can be expressed as following: *If A is true then A is true, therefore A is correct.* The proof of universality of timescale stretching based on such logic cannot be accepted.

The light curves for UBVR bands in terms of normalized flux are shown on Figure 3. The flux is normalized to the unity of flux maximum at zero redshift.

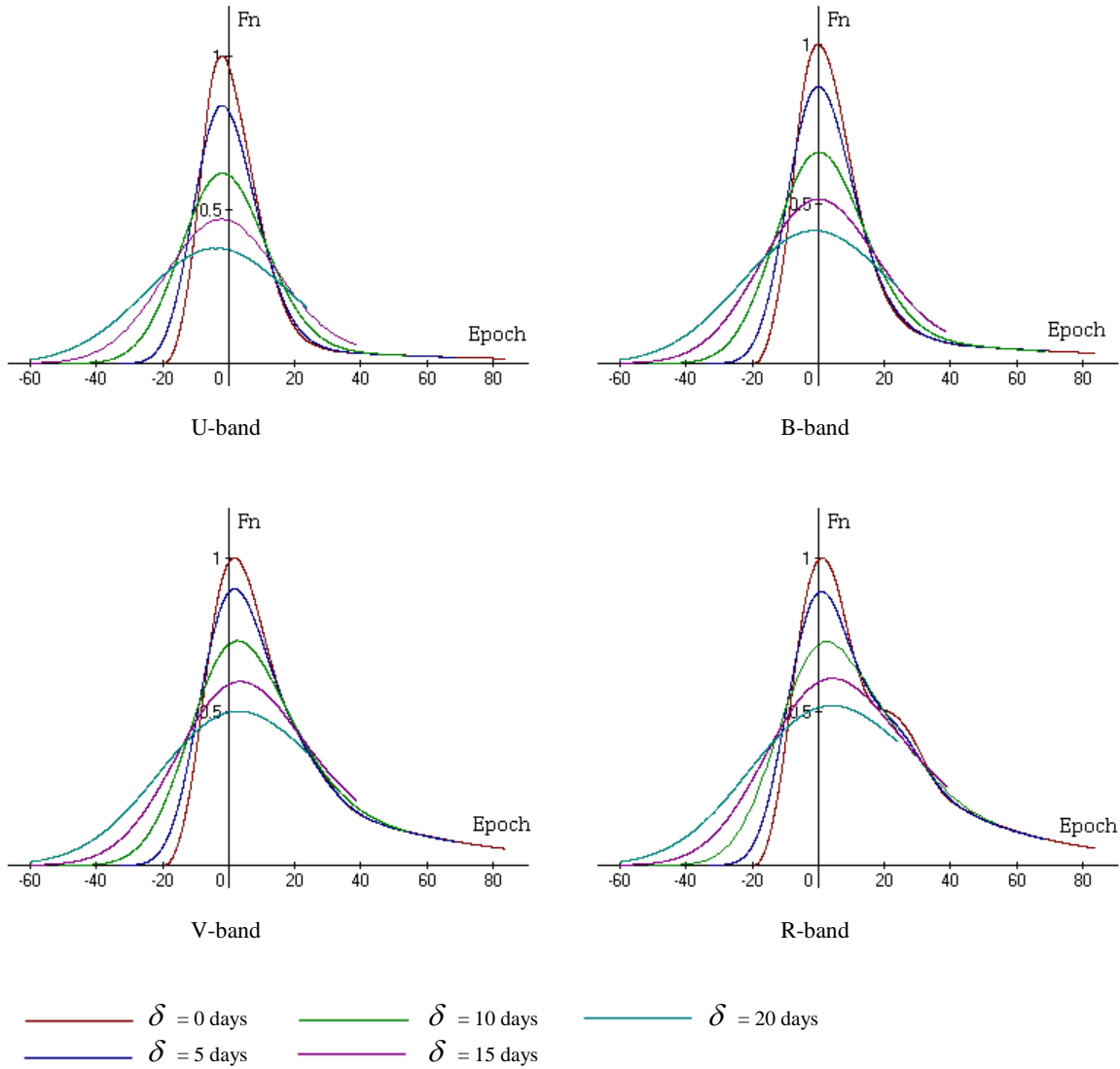


Figure 3  
Normalized flux light curves in non-shifted UBVR bands

#### 4. Light curve broadening in IYJHK bands

The magnitude light curves (non-shifted) for various standard deviations in IYJHK bands are shown on Figure 4 with some magnitude offset. Width-factor for light curves in IYJHK bands cannot be measured separately for rising and falling parts as two maximums in a curve gradually transform into one with the increase of standard deviation of travelling time. So the width-factors were determined by measuring the total width. The results are presented in Table 2 and Figure 5.

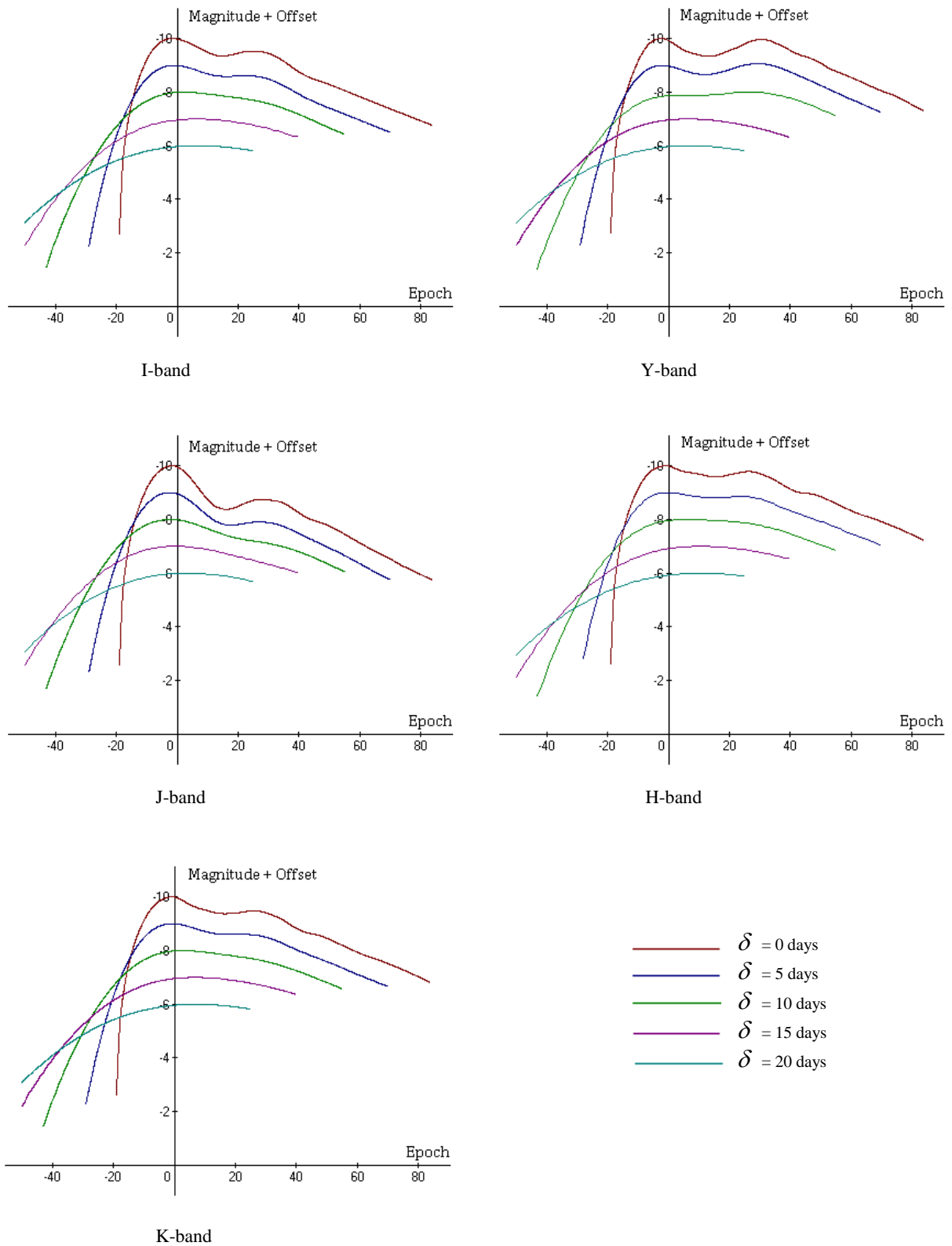


Figure 4  
The magnitude light curves in non-shifted IYJHK bands



Table 2 - Measurements of light curve broadening in UBVR bands

Fragment	Magn. Offset	I		Y		J		H		K	
		Width	w	Width	w	Width	w	Width	w	Width	w
$\sigma = 0$ days, $Z=0$											
Whole	0.5	36.2	1	49.4	1	12.9	1	42.4	1	34.0	1
Whole	1	47.7	1	61.6	1	18.8	1	54.1	1	47.9	1
Whole	1.5	57.1	1	73.3	1	49.3	1	68.8	1	61.5	1
Whole	2	70.3	1	84.8	1	57.9	1	81.8	1	72.6	1
$\sigma = 5$ days, $Z=0$											
Whole	0.5	39.9	1.10	55.6	1.13	16.9	1.31	46.10	1.09	40.0	1.18
Whole	1	52.1	1.09	68.3	1.11	25.3	1.35	61.20	1.13	53.9	1.13
Whole	1.5	63.4	1.11	94.1	1.28	55.3	1.12	75.00	1.09	67.2	1.09
Whole	2	76.8	1.09	-	-	66.8	1.15	88.60	1.08	79.9	1.10
$\sigma = 10$ days, $Z=0$											
Whole	0.5	43.6	1.20	57.7	1.17	28.1	2.18	50.00	1.18	45.2	1.33
Whole	1	61.1	1.28	73.6	1.19	53.2	2.83	68.00	1.26	62.9	1.31
Whole	1.5	75.3	1.32	-	-	69.1	1.40	-	-	77.5	1.26
Whole	2	-	-	-	-	81.6	1.41	-	-	-	-
$\sigma = 15$ days, $Z=0$											
Whole	0.5	49.2	1.36	59.5	1.20	42.3	3.28	52.8	1.25	49.8	1.46
Whole	1	-	-	-	-	65	3.46	-	-	71.4	1.49

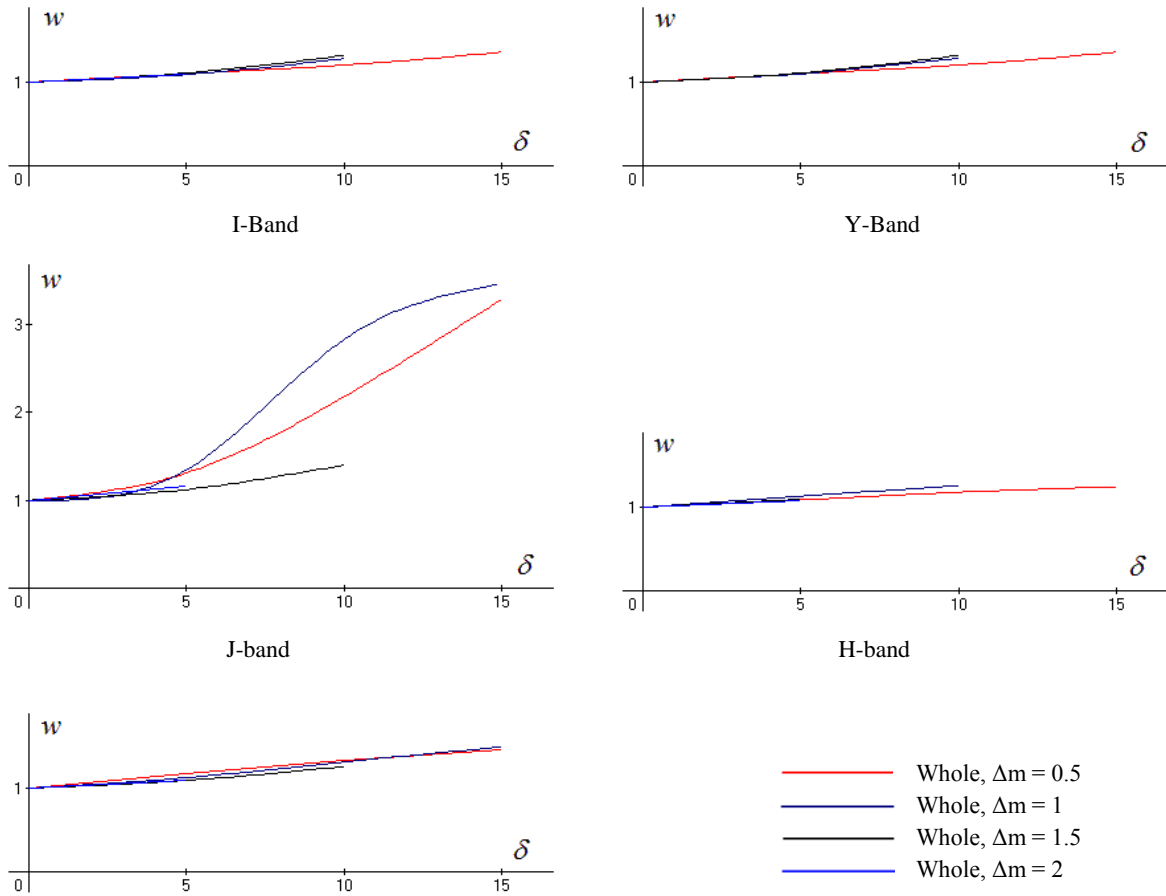


Figure 5

Dependency of the width-factor on standard deviation of photons travelling time in non-shifted IYJHK bands

The light curves for IYJHK bands in terms of normalized flux are shown on Figure 6. The flux is normalized to the unity of flux maximum at zero redshift.

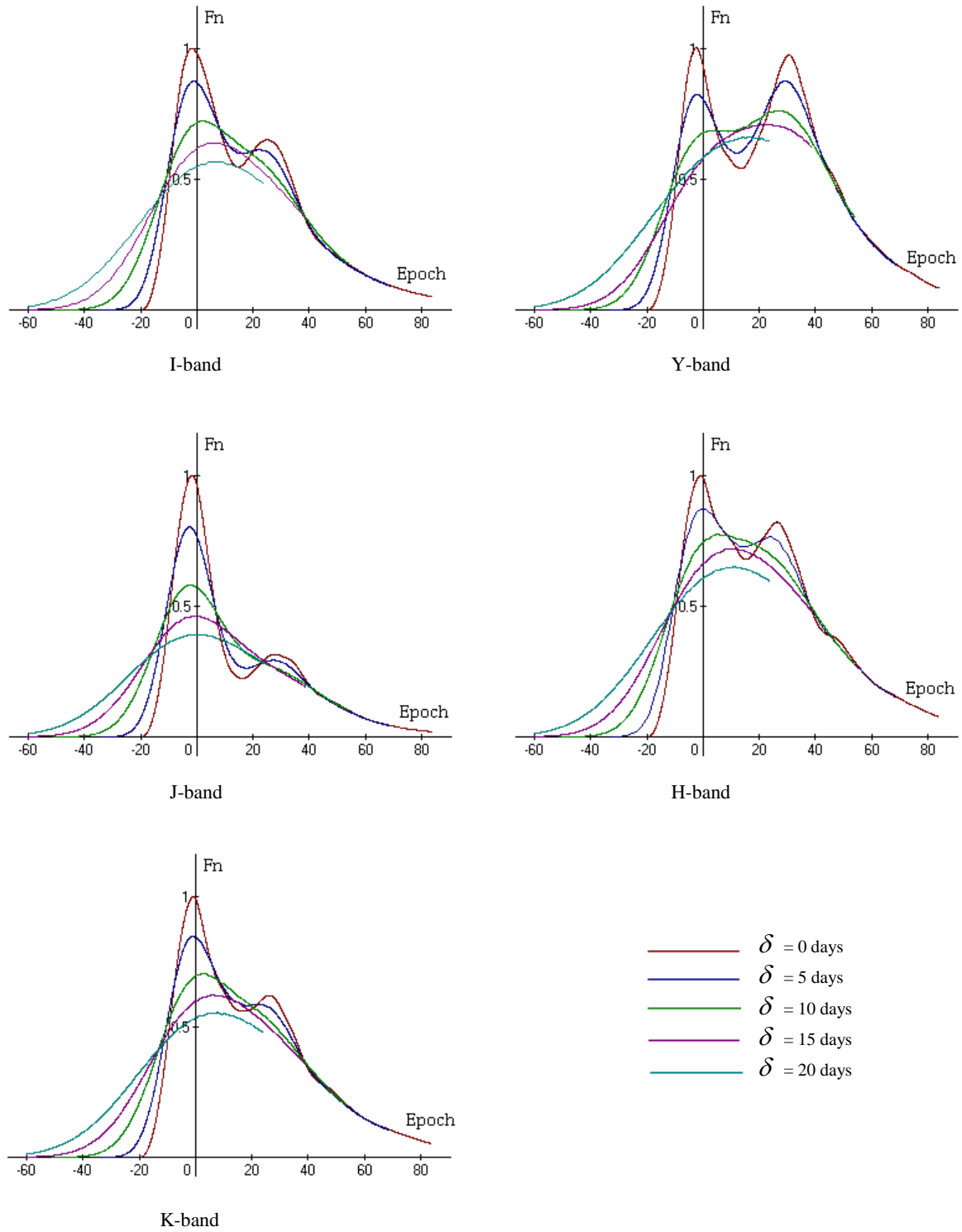


Figure 6  
Normalized flux light curves in non-shifted IYJHK bands

Figures 4 and 6 show that with the increase of the standard deviation of travelling time, the second maximum gradually disappears from the IYJKH light curves. This can be used as a good test for the validity of “tired light” paradigm. In order to validate this result against available observations I reviewed photometric measurements in non-shifted (“at rest”) I-band for 21 high-Z supernova which are listed in Table 3. These measurements are available at The Carnegie Supernova Project web-site (CSP).

Table 3 - Reviewed SN from Carnegie Supernova Project

SN	Z	Source	SN	Z	Source	SN	Z	Source
SN6699	0.311	SDSS	05D2dw	0.417	SNLS	05D1dn	0.566	SNLS
SN4679	0.332	SDSS	04D1rh	0.435	SNLS	03D4gl	0.571	SNLS
d149wcc4_11	0.342	Essence	e108wdd8_4	0.469	Essence	04D1oh	0.590	SNLS
05D2mp	0.354	SNLS	05D2bv	0.474	SNLS	04D2an	0.620	SNLS
05D4fo	0.373	SNLS	05D1ix	0.490	SNLS	04D1sk	0.663	SNLS
05D4cw	0.375	SNLS	04D1pg	0.515	SNLS	05D2bt	0.679	SNLS
SN5183	0.390	SDSS	05D2eb	0.534	SNLS	05D2ck	0.698	SNLS

Most of the SNs have the photometric measurements that do not provide clear indication of presence or absence of the second maximum in non-shifted I-band. One may consider SN 05D4cw as fitting to light curve template with two maximums but on the other hand it has only three photometric points and also fits quite well to light curve with one maximum as it is shown in Figure 7 (highlighted by red). Also it is not really known what standard deviation  $\delta$  corresponds to  $Z=0.375$  in I-band. Another striking case is 05D2bv which fits better to one-maximum curve than to two-maximum (Figure 8). The original images for Figures 7 and 8 are obtained from CSP web-site.

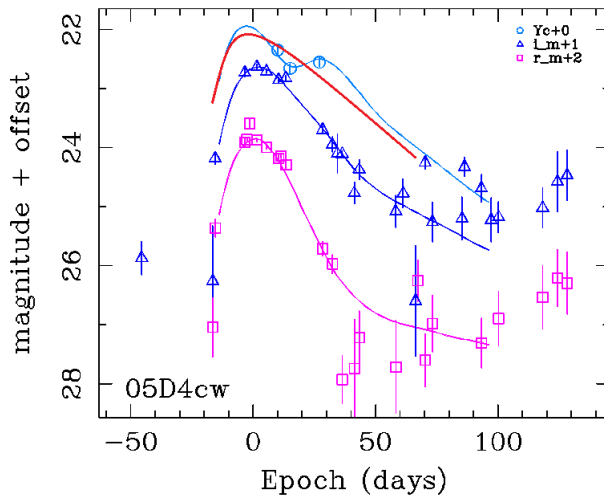


Figure 7

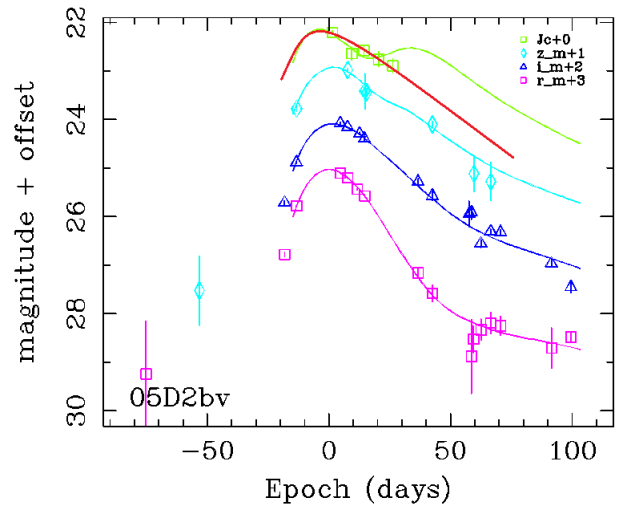


Figure 8

Table 2 and Figure 5 indicate that time-scale stretching in non-shifted IYJHK bands is generally less than in UBV bands. This can be yet another test for the validity of “tired light” hypothesis.

The photometric observations in infrared and near-infrared bands available today for high-Z supernovae cannot be used to confirm or rule out “tired light” hypothesis. These measurements should be more systematic and have better frequency and

accuracy. However it is possible to say that “tired light” hypothesis at least does not contradict to available observations in infrared and near-infrared bands.

## 5. Spectra Aging

In the presented “tired light” model the photons issued at the same time will arrive on different days. This means that the observed spectra will differ from original. The conducted simulation study shows that with the increase of standard deviation of travelling time the spectrum gets more blurry but still remains recognizable as Type Ia. Figure 9 shows spectrum at Epoch 0 for various values of  $\delta$ .

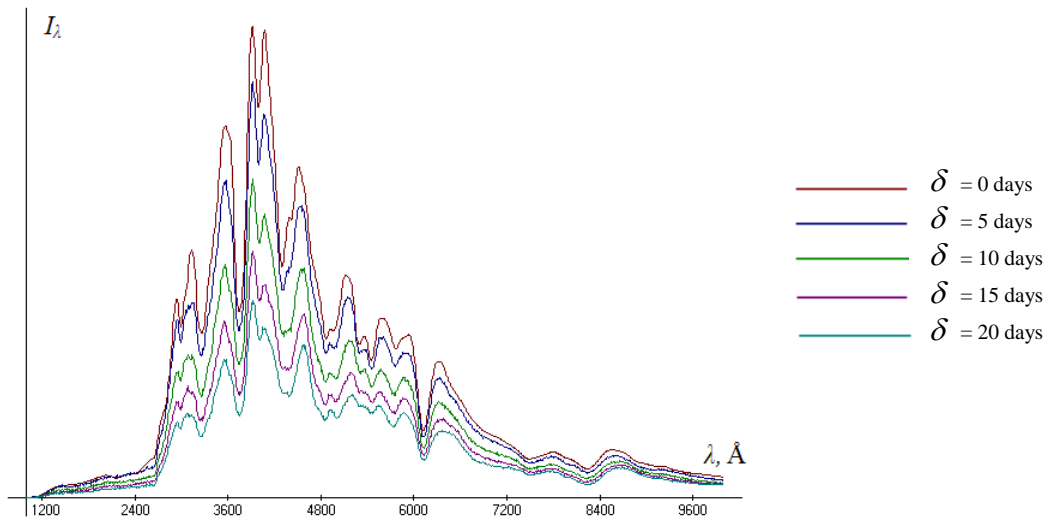
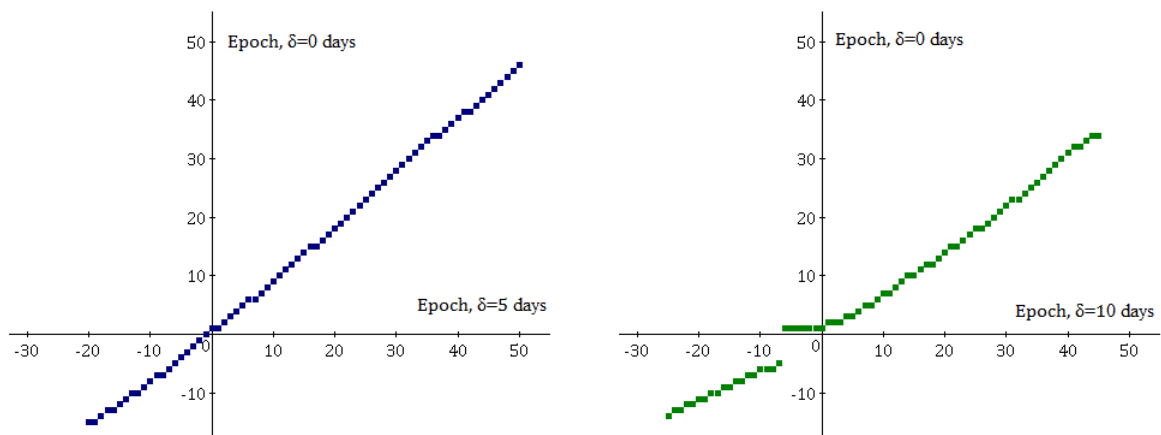


Figure 9

The fact that the spectra remain recognizable makes it possible to conduct spectra aging study using the approach suggested by Blondin et. al (2008). In order to do that I have normalized each epoch spectrum of certain standard deviation to the unity of maximum spectral flux density at this standard deviation. The wavelength range for the spectra is from 3000Å to 7000Å. Then I matched normalized spectra of each epoch at  $\delta > 0$  to normalized spectra of each epoch at  $\delta = 0$  using the least square method. The results of the matching are shown on Figure 10. The software that I have developed for matching spectra is available for download as open source and freeware (Holushko).



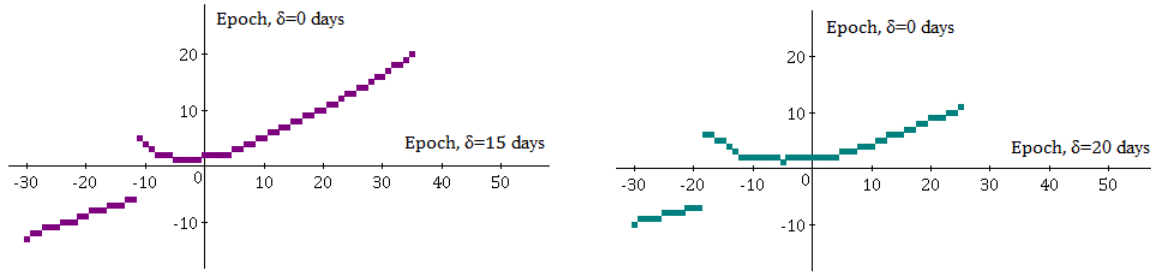


Figure 10  
Matching spectra between epochs

Blondin et al. matched spectra for epochs in the range from 0 to 30. The results of matching the simulated spectra after epoch 0 are shown on Figure 11. The matching points very well fit to right lines, which indicates that the “tired light” model is in consistency with the results of Blondin et al.

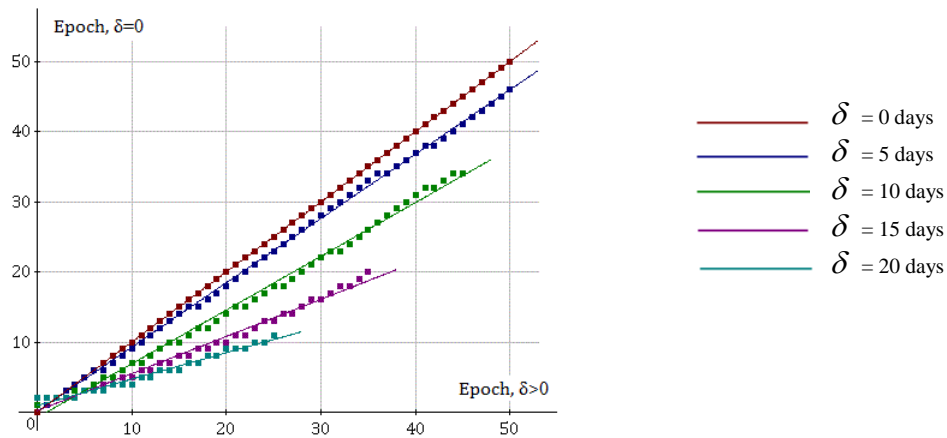
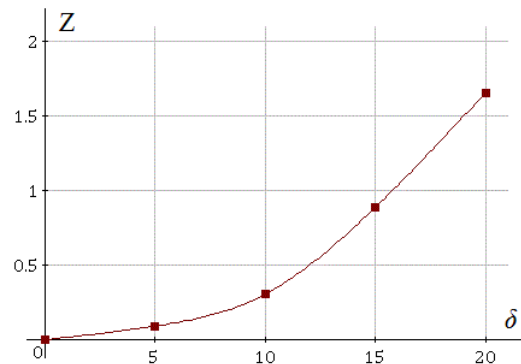


Figure 11  
Matching spectra after epoch 0

The significance of the results of spectra aging study conducted by Blondin et al. is that they were obtained without applying any subjective fitting. According to Blondin et al. the linear coefficient  $k$  of matching line is related to redshift as  $k=1/(1+Z)$ . Assuming that the results of Blondin et al. are correct it is possible to establish relation between  $Z$  and  $\delta$ . Strictly speaking it is not known if  $\delta$  is the same for all wavelengths. So in this relation  $\delta$  should be regarded as weighted average for the range 3000...7000Å. The relation between  $Z$  and  $\delta$  is shown in Table 4 and Figure 12.

Table 4

$\delta$	$k$	$Z$
0	1.000	0.000
5	0.917	0.090
10	0.766	0.305
15	0.529	0.890
20	0.376	1.657



## 6. Discussion

Redefined “tired light” model suggests the mechanism explaining “timescale stretching” effects for Type Ia supernovae. The conducted simulation study shows that observations related to light curve broadening as well as for spectra aging do not contradict to “tire light” model. Therefore ruling out “tired light” model is not justified. “Tired light” has the following advantages over “expanding Universe” model.

- It returns physical meaning to light as waves in media
- Eliminates need for mysterious “dark energy”
- Eliminates absurdity of superluminal velocity (relative to Earth) for the stars with redshift  $Z > 1$
- It is aligned with Tolman test which shows that Universe does not expand (Lerner ????, Soares ?????)

In order to confirm or rule out “tired light” or “standard cosmology” model by timescale stretching effects for Type Ia supernovae, the directions of further research should be the following:

- a) Detecting second maximum in infrared and near-infrared bands for high-Z supernovae
- b) Testing universality of stretching parameter for different bands, fragments of light curve and magnitude offsets without fitting to universally stretched light curve templates
- c) Conducting spectra aging study for epochs before 0
- d) Expanding spectra template for  $Z=0$  at least up to 120 epochs after 0 which will extend simulation results for “tired light” model

## Acknowledgements

## References